Early Journal Content on JSTOR, Free to Anyone in the World

This article is one of nearly 500,000 scholarly works digitized and made freely available to everyone in the world by JSTOR.

Known as the Early Journal Content, this set of works include research articles, news, letters, and other writings published in more than 200 of the oldest leading academic journals. The works date from the mid-seventeenth to the early twentieth centuries.

We encourage people to read and share the Early Journal Content openly and to tell others that this resource exists. People may post this content online or redistribute in any way for non-commercial purposes.

Read more about Early Journal Content at http://about.jstor.org/participate-jstor/individuals/early-journal-content.

JSTOR is a digital library of academic journals, books, and primary source objects. JSTOR helps people discover, use, and build upon a wide range of content through a powerful research and teaching platform, and preserves this content for future generations. JSTOR is part of ITHAKA, a not-for-profit organization that also includes Ithaka S+R and Portico. For more information about JSTOR, please contact support@jstor.org.

of 248 stars in the region around B. D. +35°4013. The choice of the region of the sky covered by their plates had been determined by the fact that it contained three Wolf-Rayet stars. These three stars were also put on the Mount Wilson program and the parallaxes of two of them have been finished. The results are:

B.D. +35°4001	πrel. = +0″.011 ±0″.005	16 exposures
B.D. +35°4013	πrel. = +0″.002 ±0″.003	14 exposures

The measures and reductions were made in practically the same way as that described in *Contributions from the Mount Wilson Solar Observatory,* No. 111. To derive absolute parallaxes the values found should be corrected by about +0″.002. The absolute magnitudes are then +4.1 ±0.8 and +0.9 ±1.8.

The relative parallaxes of these two stars are given by Kapteyn as +0″.004 and +0″.032, respectively, corresponding to absolute parallaxes of +0″.008 and +0″.036, while the probable errors are 0″.025 and 0″.020. However, from a comparison of the mean parallaxes as given in *Groningen* 1, page 72, with those derived by Kapteyn in *Groningen* 8, page 28, it seems likely that these parallaxes need a magnitude correction. From a curve constructed from the values given in *Groningen* 1, page 72, it would seem that the necessary corrections for the two stars are —0″.014 and —0″.025, respectively, leaving as absolute parallaxes —0″.006 and +0″.011.

It is to be regretted that it has been impossible to derive also the parallax of the third Wolf-Rayet star, B. D. +36°3956, but this will have to wait until next year.

I should like to correct here an error made in the printing of the proper motion of 61 *Cygni,* as given in this JOURNAL, February, 1916:

61_1 Cygni μ_α should read = +4″.169 instead of +4″.107.

A. VAN MAANEN.

MAGNITUDES AND COLOR OF BARNARD'S PROPER-MOTION STAR, $17^h\ 53^m\ 44^s + 4°\ 27'.8$ (1916)

The photographic and photovisual magnitudes of Barnard's large proper-motion star have been determined by direct comparisons of the field with the North Polar Standards. Exposures on Seed 27 and Cramer Isochromatic plates (the latter behind a yellow filter) were made with the 60-inch reflector of the Solar Observatory in the usual manner. Measures of two pairs of photographs exposed on different nights gave the first of the following series of results:

	Polar Comparisons	Adopted
Photographic Mag.	11.67	11.43
Photovisual Mag.	9.67	9.67
Color Index	+2.00	+1.76

The color index was also derived by the method of exposure ratios. Four isochromatic plates were exposed as follows (two nights):

With yellow filter	16, 32^s	(Yellow images)
Without filter	2, 4, 8, 16^s	(Blue images)
With yellow filter	16, 32^s	(Yellow images)

From the images on these plates it was found that the ratio of exposure times producing blue and yellow images of the same size is 0.434, which, for the emulsion used, corresponds to a color index of 1.76 ± 0.03 (P. E.) mags. This agrees well with the spectrum, which is Mb, and is much more reliable than the value given above, as the photographic magnitudes obtained from the polar comparison plates are discordant and uncertain. Since the photovisual magnitudes are probably reliable, the adopted photographic magnitude is that obtained by combining the mean photovisual magnitude with the color index.

The following are approximate results for the neighboring bright star B. D. +4°3560 ($8^m.7$):

Photographic Mag.	9.1
Photovisual Mag.	8.9
Color Index	+0.2

FREDERICK H. SEARES.

THE PHOTOGRAPHIC MAGNITUDE OF THE NINTH SATELLITE OF JUPITER.

The brightness of the Ninth Satellite of *Jupiter* has been determined by direct estimates of images on two photographs made with the 60-inch reflector. Both plates have equally-timed exposures on the Polar Standards and on the satellite